

Variability of solar/stellar activity and magnetic field and its influence on planetary atmosphere evolution

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It is shown that the evolution of planetary atmospheres can only be understood if one recognizes the fact that the radiation and particle environment of the Sun or a planet's host star were not always on the same level as at present. New insights and the latest observations and research regarding the evolution of the solar radiation, plasma environment and solar/stellar magnetic field derived from the observations of solar proxies with different ages will be given. We show that the extreme radiation and plasma environments of the young Sun/stars have important implications for the evolution of planetary atmospheres and may be responsible for the fact that planets with low gravity like early Mars most likely never build up a dense atmosphere during the first few 100 Myr after their origin. Finally we present an innovative new idea on how hydrogen clouds and energetic neutral atom (ENA) observations around transiting Earth-like exoplanets by space observatories such as the WSO-UV, can be used for validating the addressed atmospheric evolution studies. Such observations would enhance our understanding on the impact on the activity of the young Sun on the early atmospheres of Venus, Earth, Mars and other Solar System bodies as well as exoplanets.

Key words: Young Sun, solar activity, solar proxies, solar wind, atmosphere evolution, ENAs, WSO-UV, PLATO.

1. Introduction

The Sun has always played a major role in solar-planetary relations of planetary atmospheres. During the past decade the study of the history of our Sun and its influence on the evolution of Solar System planets has become an interdisciplinary effort between stellar astronomy, astrophysics of star and planet formation, astrochemistry, solar physics, geophysics, planetology, as well as climatological science. The great importance of understanding the historical evolution of our host star is obvious; the Sun's radiative energy, the plasma outflow, and various forms of transient phenomena such as shock waves, high-energy particle events during flares, and Coronal Mass Ejections (CMEs) are key factors in the formation and atmospheric evolution of the plan-

ets including Earth's biosphere (e.g., Lundin *et al.*, 2007; Lammer *et al.*, 2008, 2009).

The increase in solar luminosity over time and its climatological effect on Venus, Earth and Mars have been discussed by various authors in the past (e.g., Sagan and Mullen, 1972; Walker, 1975; Newman and Rood, 1977; Owen, 1979; Chassefière, 1996; Guinan and Ribas, 2002; Lammer *et al.*, 2003; Ribas *et al.*, 2005; Tian *et al.*, 2005; Kulikov *et al.*, 2006, 2007; Cnossen *et al.*, 2007). The early Sun, with a luminosity of about ~70% of that today, should have been too faint to prevent the early Earth and Mars from freezing. However, paleoclimate studies show that the young Earth always had liquid water and was therefore, typically warmer than during recent times and possibly heated by greenhouse gases such as CO₂, CH₄, and H₂O. Similar for Mars we know that our outer neighbor planet has traces of liquid H₂O-flows on its surface which indicates that there was once a time when Mars had a denser atmosphere and a

warmer climate.

Besides the impact of the Sun on the planetary climate, solar radiation and particle fluxes play a tremendous role in the heating of the upper planetary atmospheres. The studies based on previous observation indicated that contrary to a weaker total luminosity the young Sun was a far stronger source of energetic particles and electromagnetic radiation (e.g., Newkirk Jr., 1980; Skumanich and Eddy, 1981; Zahnle and Walker, 1982). These early studies are now more and more supported by a large number of multiwavelength (X-ray, SXR, EUV, FUV, UV, optical) observations of solar proxies which are a solid evidence that the early Sun was a much more active star than it is at present (e.g., Ribas *et al.*, 2005).

If the young Sun was more active than today, one has to study the consequences this might have had for the development of the planetary environments, including the formation and evolution of planetary atmospheres and initial H₂O inventories. To understand the impacts of the radiation and particle environment of the young Sun or other stars on early planets and their potential habitability one has to focus on the following two main subjects:

- Young stars rotate more rapidly and have an internal magnetic dynamo that is more efficient than that of the present-day Sun. This results in stronger surface magnetic fields and/or higher surface magnetic filling factors which induce enhanced “activity” in all its variations, from larger surface spots to stronger, extended solar wind.
- The optical and infrared emissions of the young star largely control the planetary atmosphere and its related climate. The X-ray, SXR, EUV photons as well as the interaction with high-energy particles and the solar/stellar wind lead to atmospheric modifications which can be observed by satellites.

In this article we focus on the evolution of the solar/stellar radiation, plasma and magnetic environment during the Sun’s (stars) life and its response on the upper atmospheres of terrestrial planets.

In Section 2 we discuss the latest research within the “Sun in time” project including uncertainties in stellar rotation periods, ages, radiation and plasma flow. Because the radiation and plasma history of our Sun can be separated into a period of a more moderate decrease from present to about 4 Gyr ago and a very early extreme period of about 500 Myr after the young Sun arrived at the Zero-Age-Main-Sequence (ZAMS) we discuss its related radiation and plasma environment separately in Section 3 and Section 4. In Section 5 we discuss the latest theoretical findings on the impact of the extreme space weather of the young Sun on the evolution of the atmospheres of Venus, Earth and Mars and the consequences for terrestrial exoplanets orbiting stars with lower masses than the Sun. Finally we discuss in Section 6 a new innovative idea how one can test the atmospheric evolution theories addressed in Section 5 by transit observations of Earth-size exoplanets, which are exposed to higher SXR/EUV fluxes of dwarf stars.

2. From “The Sun in Time” to “Stars in Time”

In our understanding of the past magnetic evolution of the Sun, solar analogs or the so-called solar proxies, Sun-like stars with different ages and activity levels play an important role. For stars with masses $\leq 1.5M_{\text{Sun}}$ and ages of about a few 100 Myr, angular momentum loss by a stellar wind brakes rotation in a way that correlates very well with the stellar age (e.g., Soderblom, 2010). The only independent variable, the stellar age, then determines the rotation period and, through the internal magnetic dynamo, magnetic activity at all levels of the stellar atmosphere, probably including characteristics of the stellar wind (Wood *et al.*, 2002, 2005; Güdel, 2007).

The age determination of a star is one of the most complex issues in stellar astrophysics. Besides the age-rotation relation stellar ages can be inferred from memberships in moving groups or stellar clusters, isochrones, chromospheric activity (CaII HK indices), and astroseismology (e.g., Soderblom, 2010). Therefore, at best, estimates are coming from a variety of indicators that are used over different timescales. In this sense, it is very hard to associate error bars since systematic errors and not random errors are the dominating factors. Specifically for the rotation periods, their relative uncertainty is proportional to the period itself since longer periods mean less active stars and are thus harder to monitor because of the smaller modulation amplitude. For the “Sun in Time” stars shown in Table 1 rotation periods are known to better than 5% in all cases. For $\kappa^1\text{Cet}$, for example, the rotation period which was determined by photometry with the MOST satellite has been determined to better than 1% (Walker *et al.*, 2007). However, for an inactive star like βHyi , the determination is significantly poorer (around 10%) (Walker *et al.*, 2007).

The observation of and detailed research on a series of near-solar-mass stars back to ages close to the ZAMS-age is therefore sufficient to reconstruct the history of our Sun and the interaction of its magnetic activity with its overall environment. This is the main goal of the so-called “Sun in Time” project which was established more than a decade ago by Dorren and Guinan (1994a, b). Table 1 shows a sample of very well observed and studied solar proxies and their main physical parameters.

Key studies within the “Sun in Time” project refer to the following main points:

- Photospheric (optical) and UV evolution with time to characterize level and extent of surface magnetic activity studied with IUE (Messina and Guinan, 2002, 2003; Dorren and Guinan, 1994a).
- The FUV studied with FUSE to infer transition-region magnetic activity (Guinan *et al.*, 2003).
- To obtain information on coronal activity the EUV was studied with EUVE (Ayres, 1997; Güdel *et al.*, 1997; Audard *et al.*, 1999; Ribas *et al.*, 2005).
- To obtain full coverage in coronal temperatures and characterize flares, X-rays were studied by ROSAT and ASCA (Dorren *et al.*, 1995; Güdel *et al.*, 1995a, 1997; Güdel, 1997). For spectroscopic coverage X-rays have been investigated by XMM-Newton (Telleschi *et al.*, 2005).

Table 1. Physical parameters and age-rotation relation of well observed and studied solar proxies in comparison with the Sun (after Güdel, 2007). L_X refers to a range of 0.1–2.4 keV as measured by the ROSAT satellite.

Star name	Type	Mass [M_{Sun}]	Radius [R_{Sun}]	$\log\left(\frac{L_X}{L_{\text{bol}}}\right)$	Rot. [days]	Age [Ga]
HD 129333 [EK Dra]	G1.5 V	1.06	0.95	-3.61	2.68	0.1
HD 72905 [π^1 UMa]	G1.5 V	1.03	0.95	-4.47	4.9	0.3
HD 206860 [HN Peg]	G0 V	1.06	0.99	-4.52	4.86	0.3
HD 39587 [χ^1 Ori]	G1 V	1.01	0.96	-4.65	5.24	0.3
HD 1835 [BE Cet]	G2 V	0.99	1.02	-4.46	7.65	0.6
HD 20630 [κ^1 Cet]	G5 V	1.02	0.93	-4.73	9.21	0.65
HD 114710 [β Com]	G0 V	1.08	0.925	-5.52	12	1.6
HD 190406 [15 Sge]	G5 V	1.01	1.1	-5.64	13.5	1.9
Sun	G2 V	1	1	-6.29	25.4	4.6
HD 146233 [18 Sco]	G2 V	1.01	1.03	—	23	4.9
HD 128620 [α Cen A]	G2 V	1.1	1.22	-6.67	~30	5–6
HD 2151 [β Hyi]	G2 IV	1.1	1.9	-6.41	~28	6.7
HD 186408 [16 Cyg A]	G1.5 V	1	1.16	—	~35	8.5

- To measure the production of non-thermal particles in magnetic coronae the radio flux has been studied with the VLA (Güdel *et al.*, 1995a, b, 1997; Gaidos *et al.*, 2000; Güdel and Gaidos, 2001).
- Magnetic cycles of solar proxies (Messina and Guinan, 2002) are measured in the optical or in X-rays. Cycles were also studied for α Cen A a G2V star (Robrade *et al.*, 2005; Ayres, 2009) and HD 81809 (G2V+G9V) (Favata *et al.*, 2004, 2008) stars.
- Differential rotation of the solar proxies as seen in stars spot measurements were studied by Messina and Guinan (2003).

The much stronger magnetic activity of the young Sun resulted in increased flaring activity, CMEs and denser and more energetic solar wind and particle fluxes. For the study of atmospheric evolution and related thermal and non-thermal escape processes the evolution of the X-ray, SXR and EUV flux is of particular interest, because it can ionize atmospheric species and dissociate molecules, that would otherwise be relatively inert, thereby initiating photochemistry that can alter an atmosphere's bulk composition as well as creating a class of radiation-dependent trace chemicals. Furthermore as a result of these photochemical processes, this wavelength range controls the temperature structure and expansion of the upper atmosphere and should have played an important role for large-scale atmospheric escape and isotopic fractionation on early Venus, Earth, and Mars.

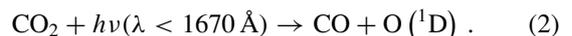
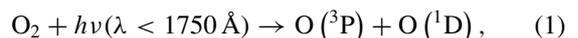
Because lower-mass stars are very common and hence may host rocky planets within their habitable zones, progress in this direction has started by expanding the “Sun in Time” project to time sequences of the high-energy emissions, CME activity, stellar winds, and flare activity to lower-mass K and M-type stars. These stars are far more numerous than the solar type stars, have long main-sequence lifetimes, and have become prime targets for terrestrial planet searches in recent years. Because of the low luminosities, their habitable zones (e.g., Kasting *et al.*, 1993; Lammer *et al.*, 2009) can be quite close to the host stars. Because these dwarf stars have deeper outer convective zones than Sun-like stars and thus possess very efficient magnetic dynamos, the expected enhanced X-ray,

SXR, EUV radiation environment combined with extreme space weather effects should play a major role in the development of the atmospheres and habitability on such planets (Scalo *et al.*, 2007; Lammer *et al.*, 2009).

3. The Moderate Activity Period: Solar Radiation, Plasma and Magnetic Environment From the Present to 4 Gyr Ago

Ultraviolet emission from the Sun and late-type stars originates in magnetic chromospheric and transition-zone regions which have been heated to temperatures of $\sim 10^4$ – 10^5 K. Based on spectral measurements obtained by IUE, Dorren and Guinan (1994a) have studied the evolution of ultraviolet (UV) line fluxes in detail and Güdel *et al.* (1997) and Ribas *et al.* (2005) extended this research to X-rays, SXR and EUV. The relevant wavelengths for thermospheric heating are the “ionizing” ones ≤ 1000 Å (e.g., Hunten, 1993), which contain only a small fraction of the present solar spectral power. Planetary atmospheres are opaque to EUV radiation (100–920 Å), as in the case of Earth, due to photo-absorption by N_2 , O_2 and O , so that the optical depth of unity at 100 Å occurs at an altitude of about 130 km.

Molecular constituents such as molecular hydrogen, carbon dioxide, molecular oxygen and N_2 (Bakalian, 2006) can also be dissociated by relatively long wavelengths in the “Schumann-Runge” range of 1750–1300 Å, according to



The solar SXR and EUV radiation with an intensity of ≥ 2 erg cm^{-2} sec^{-1} represents the predominant heat source for Earth's upper atmosphere. Dissociation can also take place below the true dissociation limit due to excitation of molecules into a state which dissociates.

Dissociation of molecular ions and ion-neutral reactions produce so-called “hot” atoms which can reach the upper thermosphere and exobase where they form planetary coronae which can be ionized and picked up by the solar wind plasma flow (e.g. Nagy *et al.*, 1981; Lammer and Bauer, 1991; Kim *et al.*, 1998; Hodges, 2000; Krestyanikova and

Table 2. Solar photon fluxes for moderate solar activity at 1 AU (Bauer and Lammer, 2004).

λ [Å]	Φ_{∞} [10^9 ph cm $^{-2}$ sec $^{-1}$]	I_{∞} [erg cm $^{-2}$ sec $^{-1}$]
1215.7 (Ly α)	300	5
1027–911	11.61	0.23
1025.7 (Ly β)	3.5	0.067
977 (C III)	4.4	0.090
911–800	8.3	0.20
800–630	2.4	0.064
630–460	4.7	0.17
584.3 (He I)	0.9	0.03
460–370	0.63	0.03
370–270	10.3	0.65
303.8 (He II)	5.4	0.35
270–205	4.5	0.36
205–153	4.6	0.49
153–100	0.4	0.06
120–80	0.3	0.066
80–40	0.33	0.108

Table 3. Solar minimum and maximum values compared to the intensity flux of medium solar activity cycle for X-ray, SXR, EUV, FUV and Lyman-alpha wavelength ranges.

Wavelength	High solar activity	Low solar activity
X-ray [1–20 Å]	~2.6	~0.4
SXR [20–100 Å]	~1.8	~0.7
EUV [100–912 Å]	~1.5	~0.8
FUV [912–1180 Å]	~1.3	~0.8
Lyman- α [1200–1300 Å]	~1.3	~0.9

Shematovich, 2006; Lammer *et al.*, 2006; Lichtenegger *et al.*, 2009; Gröller *et al.*, 2010; Valeille *et al.*, 2010).

For H₂, radiation $\lambda < 850$ Å leads to either dissociation or ionization. Atomic hydrogen possesses a strong continuum absorption cross section below 912 Å and strong Lyman-series lines, principally the intense Lyman- α line at 121.6 nm. On Earth, rocket and satellite observations have led to a detailed identification of the solar emission line spectrum; relatively accurate values for the intensities of these radiations have become available by now (e.g., Hall *et al.*, 1969; Tobiska, 1993; Tobiska and Eparvier, 1998). Solar photon fluxes at Earth for moderate solar activity are summarized in Table 2. The solar photon flux $\phi_{\infty} = \lambda I_{\infty}/(hc)$ with h the Planck constant and c the speed of light is used for the ionization process, while the SXR and EUV energy flux $I_{\infty} = (h\nu)\phi_{\infty}$ is relevant for heating of the upper atmosphere.

3.1 Short- and longtime variation of the short wavelength radiation

The solar photon fluxes vary over the 11 year solar cycle as well as over the ≈ 27 day solar rotation period owing to magnetic field inhomogeneities. Additionally to the photon flux variation, also the solar intensity fluxes differ in phase with the solar cycle. The Solar Irradiance Platform (SIP)¹, which produces full solar spectra for any observation date, can be used to obtain intensity fluxes in the SXR, EUV and FUV wavelength ranges. The results, representing quiet and

active conditions on the Sun normalized to the average flux value, are listed in Table 3.

To understand the evolution of the solar radiation beyond activity levels which can be generated during high solar activity one has to analyze solar proxies which are observed in the “Sun in Time” project. As can be seen in Table 1, the sample of solar proxies contains stars that cover most of the Sun’s main sequence lifetime from ~ 100 Myr to 8.5 Gyr. The observations, obtained with the space observatories shown in Table 4, cover a wide range between 1 Å and 8300 Å.

One should note that the actual spectra covered the wavelength range 80–400 Å in the best cases but some stellar spectra were still noisy well below 400 Å, due to the high opacity of the interstellar medium at these wavelengths. Therefore, the flux in the EUV range (100–912 Å) given in Tables 5 and 7 was partly interpolated. The approach is explained in detail in Ribas *et al.* (2005). A promising alternative to the interpolation was applied by Sanz-Forcada *et al.* (2011) to a sample of stars with exoplanets: the X-rays and UV spectral information is used to calculate a model of the corona and transition region, the stellar layers where the XUV emission originates; this model is then used to synthesize the spectral energy distribution in the whole XUV range. The application of this approach to the two solar ancestors EK Dra and κ Cet shows less than 0.15 dex of disagreement with the EUV flux calculated by Ribas *et al.* (2005).

It should be mentioned that the present solar activity

¹<http://www.spacewx.com/solar2000.html>

Table 4. Space observatories and instruments used for the relevant observations within the “Sun in Time” project (Güdel, 2007).

Space observatory	λ [Å]	E [keV]
XMM-Newton [CCDs]	0.83–83	0.15–15
XMM-Newton [gratings]	6–38	0.33–2.1
ASCA [CCDs]	1.2–31	0.4–10
ROSAT	5.2–124	0.1–2.4
EUVE [gratings]	80–760	0.016–0.15
FUSE [gratings]	920–1180	—
HST [gratings]	1150–1730	—
IUE [gratings]	1150–1950	—
UBVRI [photometry]	3500–8300	—

Table 5. Solar radiation flux enhancement as function of wavelength normalized to the present solar flux value from present to 3.9 Gyr ago (Ribas *et al.*, 2005; Güdel, 2007).

Solar age [Gyr]	t b.p. [Gyr]	X-ray [1–20 Å]	SXR [20–100 Å]	EUUV [100–920 Å]	FUV [920–1180 Å]	Lyman- α [1200–1300 Å]	UV [1300–1700 Å]
4.6	0	1	1	1	1	1	1
3.2	1.4	2	1.6	1.5	1.4	1.3	—
2.6	2	3	2	1.9	1.6	1.5	—
1.9	2.7	6	3	2.7	2.1	1.9	2.4
1.1	3.5	16	6	5.1	3.4	2.8	—
0.7	3.9	37	11	8.6	5	3.9	—

also shows intermediate periodicities in the range of several months (Rieger *et al.*, 1984) and ~ 2 years (Vecchio *et al.*, 2010). These periodicities are explained in terms of a Rossby wave instability in the solar interior (Zaqarashvili *et al.*, 2010a, b). Similar periodicities have been observed in other solar like stars (Massi *et al.*, 1998). Recently, Lanza *et al.* (2009) observed short term cyclic oscillations of the total spotted area on the photosphere of CoRoT-Exo-2a, which has been suggested to be caused by Rossby waves. The Rossby wave activity was probably enhanced in the early stage of solar evolution due to rapid rotation, therefore this aspect requires future detailed study.

The total high energy output of a stellar corona depends on the available magnetic energy and is therefore a consequence of the star’s dynamo operation. Younger and more rapidly rotating stars are more X-ray luminous. A similar behavior can be observed for UV and FUV radiation; the X-ray output decreases as the star ages and its rotation period increases. Table 5 lists the short wavelength radiation fluxes of the Sun as a function of stellar age between 0.7 and 4.6 Gyr. The fluxes are divided in six wavelength intervals are normalized to the present solar values. The radiation flux-age correlation can be expressed by the following power laws (Ribas *et al.*, 2005)

$$F_{X\text{-ray}} = 2.4t^{-1.92}, \quad (3)$$

$$F_{\text{SXR}} = 4.45t^{-1.27}, \quad (4)$$

$$F_{\text{EUUV}} = 13.5t^{-1.2} + 4.56t^{-1.0}, \quad (5)$$

$$F_{\text{FUV}} = 2.53t^{-0.85}, \quad (6)$$

$$F_{\text{Ly-}\alpha} = 19.2t^{-0.72}, \quad (7)$$

where F is the flux at 1 AU in $\text{erg cm}^{-2} \text{ s}^{-1}$ and t the stellar age in Gyr. As Ribas *et al.* (2005) give no power

laws for the evolution of fluxes longward of 1200 Å, we used their equation (2) and the observational data given in their table 6 to estimate the flux evolution between 1200–1700 Å. Ribas *et al.* (2005) give observed emission line fluxes of six solar analogs (including the Sun), eight lines have wavelengths > 1200 Å. The major flux contributor in the 1200–1300 Å range is the Lyman- α line (1216 Å), but unfortunately it has only been measured for three stars. As the flux of the Si III line (1206 Å) is negligible compared to Lyman- α , we use the Lyman- α power law as a proxy for the whole 1200–1300 Å range. The range between 1300–1700 Å, which approximately covers the important “Schumann-Runge” range, contains six strong emission lines. One can see from Table 3 that the Sun during its maximum solar cycle can be used as a proxy for the Sun when it had a similar flux in average quiet activity conditions ~ 1.5 –2 Gyr ago.

3.2 The solar magnetic field

Rotation and convection lead to the solar dynamo activity, which results in complex magnetic fields in the solar interior and atmosphere. Rotation rate and convective mixing length have been continuously changing during the solar evolution, therefore the magnetic field structure in the solar atmosphere/corona was most likely different in the early stage of solar evolution from what is observed today. Young solar analogs with fast rotation ≤ 5 days show polar spots in addition to the present day low latitude Sunspots (Strassmeier, 2009). It has been suggested that a dominance of the Coriolis force over the buoyancy force in the dynamics of rising magnetic flux tubes tends to the polar appearance of magnetic flux in fast rotators (Schüssler and Solanki, 1992; Schüssler *et al.*, 1996). On the other hand, the meridional advection also may lead to the transport of photospheric magnetic flux towards the polar regions (Schrijver and Title, 2001; Mackay *et al.*, 2004).

Table 6. Stars with mass loss rates determined from Ly α absorption. Spectral types and mass loss rates taken from Wood *et al.* (2005). Ages determined via gyrochronology (t_{gyro}), chromospheric activity (t_{chromo}) and isochrones (t_{iso}) were taken from Barnes (2007). Other ages (t_{other}) are taken from the references in the last column ([1] Eggenberger *et al.* (2004); [2] Soderblom and Mayor (1993); [3] King *et al.* (2010); [4] Eggen (1996); [5] DeWarf *et al.* (2010); [6] Eggenberger *et al.* (2008); [7] Montes *et al.* (2001)).

Star	Spectral type	\dot{M} (\dot{M}_{\odot})	t_{gyro} (Gyr)	t_{chromo} (Gyr)	t_{iso} (Gyr)	t_{other} (Gyr)	Method & Ref.
α Cen	G2V+K0V	2	4.4 \pm 0.5	5.62	7.84	6.5 \pm 0.3	Astroseismology [1]
61 Vir	G5V	0.3	4.2 \pm 0.7	5.6	8.96	—	—
ξ Boo	G8V+K4V	5	0.23 \pm 0.02	0.23	<0.76	0.3	UMa moving group [2]
ϵ Ind	K5V	0.5	1.0 \pm 0.1	—	—	4 \pm 0.3	Age of BD companion [3]
61 Cyg A	K5V	0.5	2.0 \pm 0.2	2.36	<0.44	6 \pm 2	61 Cyg moving group [4]
36 Oph	K1V+K1V	15	0.59 \pm 0.07	1.14	—	2	Activity-rotation [5]
70 Oph	K0V+K5V	100	1.3 \pm 0.2	1.09	—	6.2 \pm 1	Astroseismology [6]
ϵ Eri	K1V	30	0.44 \pm 0.05	0.66	<0.6	0.65	Hyades supercluster [5]
Proxima Cen	M5.5V	<0.2	4.4 \pm 0.5	5.62	7.84	6.5 \pm 0.3	Ages adopted from α Cen
EV Lac	M3.5V	1	—	—	—	0.3	UMa moving group [7]

High-latitude appearance of Sun-spots could change the solar wind structure (possible interchange of fast and slow solar wind locations) and may also have affected the initiation latitudes of solar flares/CMEs. This could completely change the solar activity patterns, butterfly diagram and influence of solar wind/CMEs on the Solar System. However, this would have happened in the very early stages of the solar evolution, when the Sun just arrived at the ZAMS. The solar rotation period was ~ 10 days 4 Gyr ago (see Table 1), which is faster than the nowadays Sun, but still slower than at the time when the young Sun arrived at the ZAMS. This rotation rate does not affect too much on appearance of sunspot and therefore the magnetic structure of solar atmosphere/corona probably has not changed significantly during the past 4 Gyr.

3.3 The time evolution of the solar wind

As a consequence of solar/stellar activity, plasma winds are ejected from the star but their detection in solar proxies is very difficult. Although the wind formation and acceleration mechanisms are still not completely understood at the Sun, there is much evidence that magnetic fields play an important role for acceleration of the wind, for its heating, or for guiding the wind at least out to the Alfvén radius at which the stellar wind speed reaches the value of the Alfvén speed. Beyond the Alfvén radius the stellar wind is super-Alfvénic. One of the best indirect confirmation of the presence of magnetized winds in young stars is their spin-down to slower rotation (as shown in Table 1) after their arrival at the ZAMS because the outward flowing plasma carries away angular momentum from the star.

Potential detection methods of stellar winds include the measurement of thermal radio emission from the plasma flow (Lim and White, 1996; Gaidos *et al.*, 2000), and signatures of charge exchange in X-ray spectra (Wargelin and Drake, 2001). The most successful approach so far is an indirect method which makes use of the Lyman- α absorption in astrospheres due to the interaction between stellar winds and the interstellar medium (ISM) (e.g., Wood, 2004). When stellar winds collide with the ISM, at further distance from the star, a termination shock where the wind is shocked to subsonic speeds is formed. The heliosphere is permeated by interstellar H I with $T \sim (2-4) \times 10^4$ K

(Wood, 2004; Wood *et al.*, 2005). Much of this gas is piled up between the heliopause and the bow shock producing a hydrogen wall that can be detected by UV observations as an absorption signature in the Lyman- α line. The excess absorption from the Sun’s own hydrogen wall is, due to the deceleration of the ISM relative to the star, redshifted, while that of astrospheres are blueshifted. The measurable absorption depths are then compared with results from hydrodynamic model calculations (Wood *et al.*, 2005). We note that the amount of astrospheric absorption scales with the wind ram pressure, and the stellar wind velocity has to be assumed to derive the mass loss rate (Wood and Linsky, 1998).

A systematic study of all derived mass-loss rates of main sequence stars which were observed with this method indicates that dM_L/dt per unit stellar surface correlates with stellar activity and hence the stellar X-ray surface flux F_X (Wood *et al.*, 2005)

$$\frac{dM_L}{dt} \propto F_X^{1.34 \pm 0.18}. \quad (8)$$

If we extrapolate the above power law up to the observed X-ray saturation limit $F_{X,\text{sat}}$ of Sun-like stars of about 2×10^7 erg cm $^{-2}$ s $^{-1}$, we would suggest that the mass loss of plasma from the youngest stars in Table 1 is about 1000 times larger compared to the present day Sun. However, recent observations indicate that this power law relation breaks down for the most active stars for which $F_X \geq 8 \times 10^5$ erg cm $^{-2}$ s $^{-1}$ (Wood *et al.*, 2005).

Table 6 shows the stars observed by Wood *et al.* (2005) and their mass loss rates in units of the solar mass loss rate. Note that only one star from the “Sun in Time” sample in Table 1 has a measured mass loss rate (α Cen). If two spectral types are given in the second column, this indicates that two stars share the same astrosphere and the mass loss rate is a combined value for two stars. Age estimates for the stars as determined by different methods are also given. These include ages from gyrochronology (i.e. rotation-age relations), chromospheric activity, isochrones, and other methods, which include astroseismology, membership in moving groups, or known ages of binary companions. The sometimes discrepant results illustrate the difficulties in a

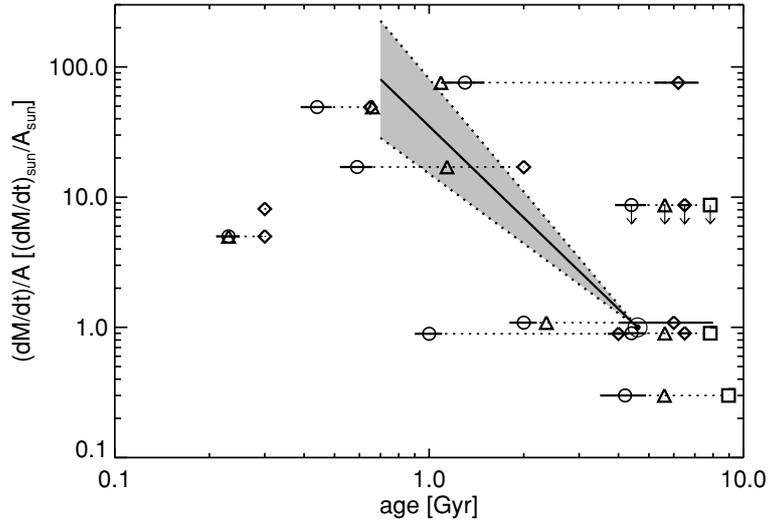


Fig. 1. Mass loss rates normalized to stellar surface area as a function of age for stars in Table 6. The adopted mass loss rate of the Sun is $2 \times 10^{-14} M_{\odot} \text{ yr}^{-1}$. The shaded area denotes the evolution $\dot{M} \propto t^{-2.33 \pm 0.55}$ derived from Eq. (3) and $F_X \propto t^{-1.74 \pm 0.34}$ (Ayres, 1997) by Wood *et al.* (2005). The different symbols (circle, triangle, square, diamond) represent different age determination methods (t_{gyro} , t_{chromo} , t_{iso} , t_{other} ; cf. Table 6).

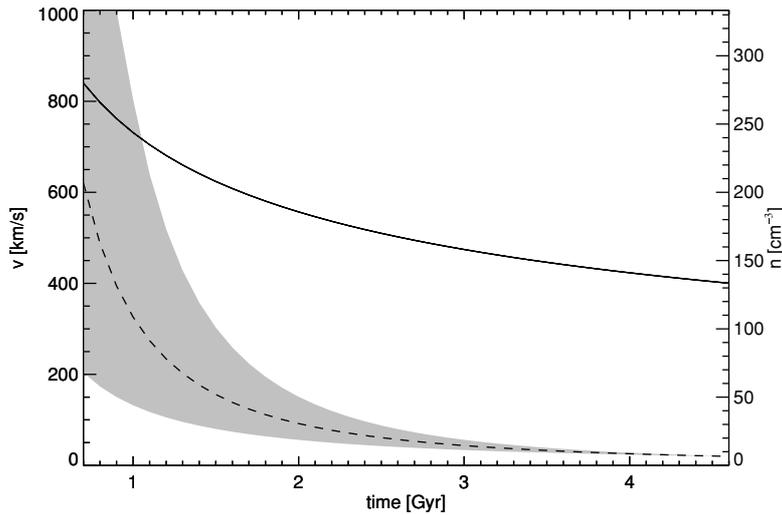


Fig. 2. Average stellar wind density (dashed line) and velocity (solid line) expected at 1 AU from present time (4.6 Gyr) to 0.7 Gyr, according to mass loss observations of Sun-like stars (Wood *et al.*, 2005) and the power laws of Newkirk (1980). The shaded area represents to uncertainties of the solar wind density in time.

reliable determination of stellar ages. Figure 1 shows the stellar mass loss rates as a function of age. The mass loss rates are normalized to stellar surface area to simplify comparison between stars with different spectral types, as well as two stars sharing the same astrosphere. Ages determined by different methods for the same star are indicated by different symbols and connected by dotted lines. The shaded area denotes the mass loss evolution derived by Wood *et al.* (2005), who used their power law fit to the X-ray surface flux (cf. Eq. (3)) and the evolution of the X-ray flux ($F_X \propto t^{-1.74 \pm 0.34}$) of Sun-like stars (Ayres, 1997) to construct the scaling $\dot{M} \propto t^{-2.33 \pm 0.55}$. As one can see from Fig. 1 stars at the lower age (i.e., high activity) limit indicate a mass loss which is about 100 times the present solar value, while the mass loss drops toward higher activity levels to about 10 times the solar value. We will discuss the

possible reason for this breakdown between X-ray activity and wind-mass loss relation in the next section in more detail.

For studying the effect on planetary environments against solar/stellar wind induced non-thermal escape, the solar/stellar wind density n_{sw} and velocity v_{sw} in the vicinity of a planet are key parameters. These parameters for a Sun-like star can be estimated as a function of the age from the observations discussed above, together with an empirical model for an age-dependence of the solar/stellar wind velocity (Newkirk Jr., 1980; Grießmeier *et al.*, 2004, 2007; Lichtenegger *et al.*, 2010)

$$v_{\text{sw}} = v_* \left(1 + \frac{t}{\tau} \right)^{-0.4}, \quad (9)$$

the minimum estimated solar/stellar wind density

$$n_{\min_{\text{sw}}}(t) = n_{\min*} \left(1 + \frac{t}{\tau}\right)^{-1.26}, \quad (10)$$

an average solar/stellar wind density

$$n_{\text{ave}_{\text{sw}}}(t) = n_{\text{ave}*} \left(1 + \frac{t}{\tau}\right)^{-1.86}, \quad (11)$$

and maximum solar/stellar wind density

$$n_{\max_{\text{sw}}}(t) = n_{\max*} \left(1 + \frac{t}{\tau}\right)^{-2.46}. \quad (12)$$

The proportionality constants (v_* , $n_{\text{ave}*}$, $n_{\min*}$, $n_{\max*}$) are determined by the present day ($t = 4.6$ Gyr) solar wind conditions at 1 AU (Earth orbit: $v_{\text{sw}} \approx 400$ km s⁻¹; $n_{\text{sw}} \approx 6.5$ cm⁻³) resulting in $n_{\min*} = 4.6 \times 10^3$ cm⁻³, $n_{\text{ave}*} = 1.04 \times 10^5$ cm⁻³, $n_{\max*} = 2.35 \times 10^6$ cm⁻³ and $v_* = 3200$ km s⁻¹ for the nominal case. The time constant $\tau = 2.56 \times 10^{-2}$ Gyr according to Newkirk Jr. (1980) and t is the time in Gyr. Figure 2 shows the expected enhancement in solar wind density and velocity as function of age from present time back to about 3.9 Gyr ago.

The density scaling (Eqs. (5)–(7)) was derived from the measured mass loss rates from Wood *et al.* (2005), hence using $\dot{M} = 4\pi r^2 m_p n_{\text{sw}} v_{\text{sw}}$, with proton mass m_p together with Eqs. (4)–(6) and $r = 1$ AU yields a similar evolution of the mass loss rate as indicated by the shaded area in Fig. 1, valid for main-sequence G-type stars with an age of at least 0.7 Gyr (Grießmeier *et al.*, 2007; Lichtenegger *et al.*, 2010).

4. The Extreme Period: The First 500 Myr After the Sun's Arrival at the ZAMS

The first 500 Myr after the origin of a terrestrial planet and the evolution of its host star to the ZAMS are the most crucial period for the development of a biosphere. During this early evolutionary stage a star (including the young Sun) is more active at short wavelengths and plasma outflows can therefore remove a planet's atmosphere or initial water inventory. To understand how the atmospheres of terrestrial planets were affected by the young Sun it is of tremendous importance to gain more knowledge from astrophysical observations of magnetic properties, location of activity centers, radiation, plasma outflow (e.g., winds, CMEs), mass loss, and rotation periods, of very young solar proxies and stars with ages <500 Myr.

4.1 The magnetic field and plasma outflow of the early young Sun

The rotation period of the young Sun was much faster (~few days) during the first 500 Myr after the arrival at the ZAMS, therefore the activity and magnetic field configuration was most likely quite different from the current appearance. Star-spots on the young solar analog EK Dra which rotates with a period of ~2.68 days are observed to be formed everywhere; from the low latitudes up to polar regions (Strassmeier, 2009). The young Sun probably had the same structure of active regions. The high latitude Sun-spots could change the structure of coronal magnetic fields as well as the formation of prominences, which might

lead to a predominant occurrence of flares and CMEs near the polar regions. The polar CMEs may propagate significantly above the ecliptical plane, which may cause more convenient conditions for planetary evolution. However, not much is known about solar magnetic dynamo during that time. It is not well constrained what the duration of activity cycles was, and also the butterfly diagram and the other laws of current solar activity just after the arrival of the Sun at the ZAMS. Therefore, the understanding of solar influence on planets requires further modeling of solar magnetic dynamos with all consequences in appearance of magnetic flux at the surface and in the possible structure of coronal magnetic field in the young Sun. This may be related to the fact that very young stars rotate fast and their active regions appear therefore at high latitudes. Furthermore, if the magnetic field becomes more akin to a global dipole, then the escape of the stellar plasma may be more inhibited (Wood *et al.*, 2005). Thus, a possibly preferred appearance of CMEs at polar regions during the fast rotating young Sun may lead to the propagation of shocks and magnetic ropes and clouds off the ecliptic plane before 4 Gyr ago, which might completely change the current picture of solar/stellar-planetary atmosphere-magnetosphere interaction processes (Zaqarashvili *et al.*, 2011). All these questions need to be answered in order to understand the influence of the young Sun on Solar System dynamics after its arrival at the ZAMS.

4.2 X-ray to UV fluxes after the Sun arrived at the ZAMS

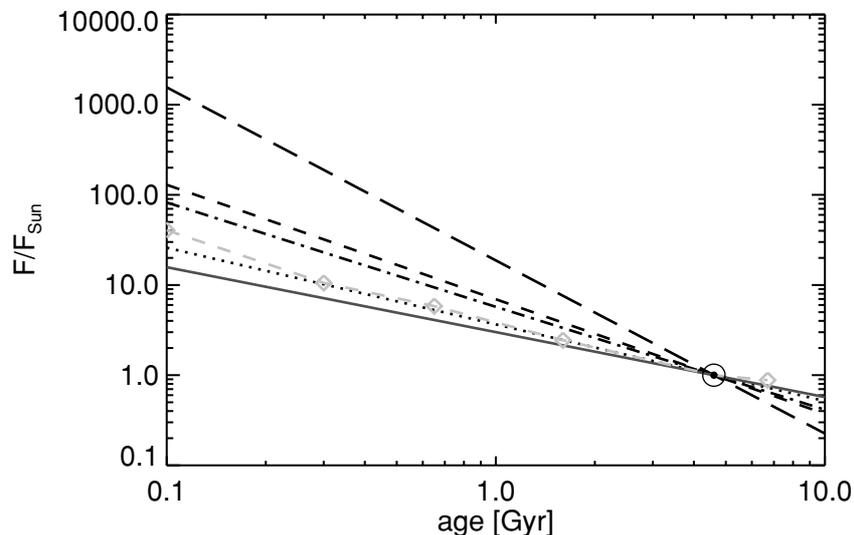
One can see from Table 1 that the “Sun in Time” sample contains 6 nearby G0 V–G5 V stars (EK Dra, π^1 UMa, HN Peg, χ^1 Ori, BE Cet, κ^1 Cet) that have known rotation periods and well-determined temperatures, luminosities, and metallicities which fall into the early age domain where $t \leq 0.65$ Gyr after the Sun's arrival at the ZAMS.

Table 7 shows similar as Table 5 integrated radiation enhancement factors according to the analyzed data for various wavelength regions as a function of a stellar or solar age between 0.1–0.7 Gyr. As in Table 5 the fluxes are divided in six wavelength intervals and are normalized to the present solar values, calculated by using the power laws as given in Section 3.1. The sums of observed emission line fluxes, are used as proxies for the evolution in the full UV range. The evolution of the integrated fluxes is displayed in Fig. 3. According to Ribas *et al.* (2005) the main uncertainties are the measurement error of the data, the calibration uncertainty of the various instruments and the intrinsic stellar variability which results in the total uncertainty of the power law slopes of ≤ 0.1 . The slope of the normalized fluxes becomes less steep with increasing wavelength, except for the 1300–1700 Å interval. By comparing the data obtained from solar proxies with ages of about 4.6–0.7 Gyr (Table 5) with those which have ages <0.7 Gyr (Table 7) one can see that the upper atmospheres of early Venus, Earth, Mars and Saturn's satellite Titan should have been tremendously affected by X-rays, SXR and EUV radiation.

Short-wavelength radiation not only gets stronger toward younger solar analogs, it also becomes spectrally harder. This property is by no means trivial. In principle, increased luminosity could be achieved by populating the stellar surface with a larger number of similar magnetic active re-

Table 7. Solar radiation flux enhancement as function of wavelength normalized to the present solar flux value during the first 500 Myr after the young Sun arrived at the ZAMS (Ribas *et al.*, 2005; Güdel, 2007).

Solar age [Gyr]	t b.p. [Gyr]	X-ray [1–20 Å]	SXR [20–100 Å]	EUV [100–920 Å]	FUV [920–1180 Å]	Lyman- α [1200–1300 Å]	UV [1300–1700 Å]
0.7	3.9	37	11	8.6	5	3.9	—
0.65	3.95	43	12	9.4	5.3	4.1	5.8
0.6	4.0	50	13	10	5.7	4.3	—
0.55	4.05	59	15	11	6.1	4.6	—
0.5	4.1	71	17	13	6.6	4.9	—
0.45	4.15	87	19	14	7.2	5.3	—
0.4	4.2	109	22	17	8	5.8	—
0.35	4.25	141	26	19	9	6.4	—
0.3	4.3	189	32	23	10	7.1	10.6
0.25	4.35	268	40	28	12	8.1	—
0.2	4.4	412	54	37	14	9.6	—
0.15	4.45	715	77	51	18	11.8	—
0.1	4.5	1558	129	82	26	15.8	40.5

Fig. 3. Evolution of the fluxes at 1 AU in different wavelength regimes (X-ray: 1–20 Å, long dashes; SXR: 20–100 Å, dashed line; EUV: 100–920 Å, dash-dotted line; FUV: 920–1180 Å, dotted line; Ly- α : 1200–1300 Å, grey solid line; UV: 1300–1700 Å, grey dashed line with diamonds) normalized to the present solar values. Power laws for the wavelength ranges were taken from Ribas *et al.* (2005). Line fluxes in the UV (1300–1700 Å) were summed up directly from observational data of strong emission lines (table 6 in Ribas *et al.*, 2005).

gions, which would not imply a change in hardness.

Observationally, spectral hardness has been characterized in terms of the average electron temperature of the emitting stellar coronal plasma. While the solar coronal plasma shows characteristic temperatures in the 1–2 MK range (e.g., Peres *et al.*, 2000), the overall trend for solar analogs with ages between 100 Myr and 5 Gyrs can be given as,

$$T_{\text{cor}} = 0.34L_X^{0.25} \quad [\text{K}] \quad (13)$$

where T_{cor} is the average coronal temperature in K and L_X is the X-ray luminosity in erg s^{-1} (Güdel *et al.*, 1997; Güdel, 2007; Telleschi *et al.*, 2005). For a young, near-ZAMS solar analog, $T_{\text{cor}} \approx 10$ MK, which results in considerably harder continuous X-ray radiation from these objects than from the present-day non-flaring Sun. This may be very important for modifying the photochemistry of early planetary atmospheres as the penetration depth is larger for high-energy

(harder) photons.

The cause of the hardening with increasing luminosity is an increasing population of coronal loops with maximum temperature around ~ 10 MK with increasing rotation of the star (e.g. Sanz-Forcada *et al.*, 2003). The nature of these loops is not clear in detail, but an attractive possibility is the operation of flares. For flares, a similar temperature-luminosity relation applies in X-rays (Güdel, 2004). Statistical studies during the past decade have addressed the problem of the occurrence rate of flares as a function of their total radiative energy (e.g., Audard *et al.*, 2000; Güdel *et al.*, 2003). The occurrence rate distributions are steep power-laws in energy. The summation of all small flares unresolved in the light curves may suffice to generate the entire observed radiation including “quiescent radiation”. Toward more active stars, progressively larger flares contribute to this baseline level; because larger flares are hotter, the summed, average radiation also becomes harder. This

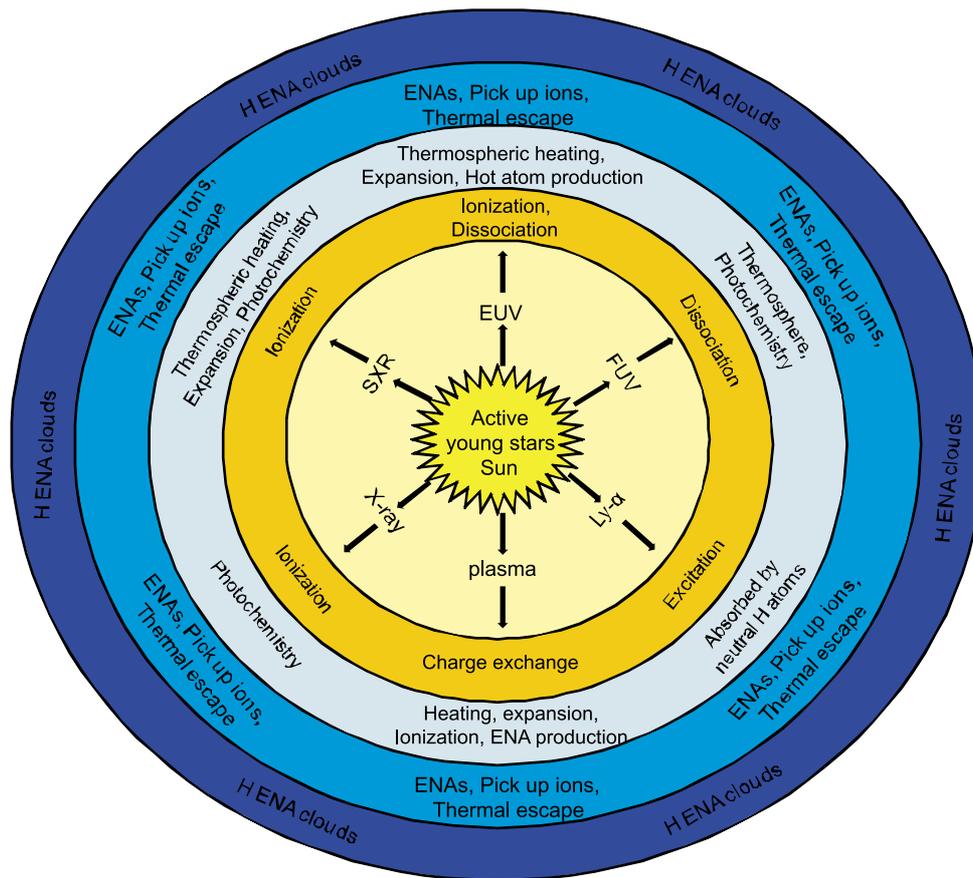


Fig. 4. Illustration of processes and response of upper atmospheres related to the radiation and plasma environment of active stars or the young Sun.

model is also supported by the *continuous* presence of non-thermal electrons in extremely active young solar analogs (Güdel, 2007) that are presumably accelerated by frequent flares. One would thus expect not only more intense X-rays in young solar analogs or around the young Sun itself, but also an increased level of flare-produced hard X-rays, gamma rays, ejected energetic particles and CMEs.

From this observational evidence one can expect that early planetary thermospheres should have been extremely hot and most likely expanded (Lammer *et al.*, 2008; references therein). In Section 5 we will discuss the problems regarding the “survival” or stability of planetary atmospheres during such extreme solar/stellar activity periods.

5. Solar Activity Response of Planetary Upper Atmospheres

In recent studies by Kulikov *et al.* (2006, 2007), Lammer *et al.* (2006, 2008), Lundin *et al.* (2007), Tian *et al.* (2008a, b) and illustrated in Fig. 4 it was shown that enhanced solar EUV radiation and plasma (winds, CMEs) should result in permanent forcing of the upper planetary atmospheres, which can ionize, heat, expand, chemically modify, and erode it during the early phase of a planetary lifetime. In some cases the thermosphere may expand beyond an atmosphere protecting magnetospheric shield. Highly irradiated terrestrial planets can thus be even in danger of being stripped of their whole atmospheres (Lundin *et al.*, 2007; Lammer *et al.*, 2009; Lichtenegger *et al.*, 2010).

When the temperature of the thermosphere is large, a significant fraction of the light constituents of the upper atmosphere attain velocities above the escape value, and the thermosphere may even become non-hydrostatic and start to expand accompanied by adiabatic cooling to several planetary radii (Chassefière, 1996; Tian *et al.*, 2005, 2008a, b, 2009; Kulikov *et al.*, 2007; Lammer *et al.*, 2008).

On the basis of their energy source, we can separate two main escape categories: thermal escape and non-thermal escape. For thermal escape, atoms at the exobase level have more or less a Maxwellian velocity distribution that is determined by the exobase temperature. Particles with velocities which are larger than the escape velocity of the planet can escape from the planet (e.g., Chamberlain, 1963). In extreme cases when the mean thermal energy of the major species at the exobase level (where the mean free path of the particles is comparable to the atmospheric scale height) exceeds their gravitational potential energy, thermal escape can be considered as “blow-off” (Öpik, 1963). Assuming hydrostatic equilibrium, Kulikov *et al.* (2007) showed that terrestrial upper atmospheres exposed to enhanced EUV radiation of the early Sun could be heated to extremely high exobase temperatures which could be over 10 000 K. Tian *et al.* (2008a, b) studied the response of thermospheres against high EUV fluxes in more detail and discovered that one can classify the thermosphere response into two thermal regimes.

In regime I, the thermosphere is in hydrostatic equilib-

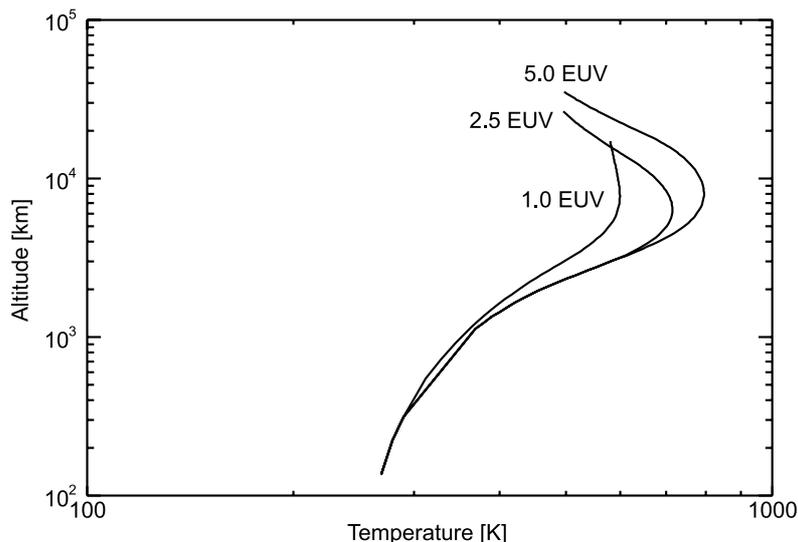


Fig. 5. Thermospheric temperature profiles for a H-rich upper atmosphere of an Earth-like planet between the mesopause and exobase level for present day EUV (1 EUV), 2.5 EUV and 5 EUV. According to Tian *et al.* (2005); for an EUV flux which is 5 times that of the present Sun the exobase moves upward to about 7 Earth radii.

rium in which the bulk atmosphere below the exobase can be considered as static. This is the situation in the thermospheres of the Solar System planets today. However, if the Sun's EUV flux or that of another star was/is much higher than that of the present Sun the thermosphere can enter a second regime.

In regime II the upper atmosphere becomes non-hydrostatic. In the hydrodynamic flow regime, the major gases in the thermosphere can expand very efficiently to several planetary radii above the planetary surface. As a result light atoms experience high thermal escape rates and expand to large distances (Tian *et al.*, 2008a, 2009) where the whole atmosphere may be eroded by solar/stellar wind plasma flow related non-thermal escape processes (Khodachenko *et al.*, 2007; Lammer *et al.*, 2007; Lundin *et al.*, 2007; Lichtenegger *et al.*, 2010).

5.1 Hydrogen-rich upper atmospheres of terrestrial planets

Terrestrial planets may evolve through phases where their upper atmospheres are hydrogen-rich. There is some evidence from planetary formation theories that early Venus once lost its initial H₂O inventory during an episode of a runaway and/or wet greenhouse effect via thermal and non-thermal atmospheric escape processes which could have been related to the form of “hydrodynamic” outflow. These extreme conditions may be reached in H-rich primordial thermospheres heated by the strong EUV flux of the young Sun (Sekiya *et al.*, 1980a, b; Watson *et al.*, 1981; Zahnle and Walker, 1982; Tian *et al.*, 2005).

Sekiya *et al.* (1980a, b) and Watson *et al.* (1981) in their early works studied hydrodynamic conditions of H-rich thermospheres from early Earth or Venus due to solar EUV heating by applying idealized hydrodynamic equations. From their thermospheric model Watson *et al.* (1981) obtained exobase levels for an Earth which is exposed to EUV fluxes of ~ 5 times that of the present Sun of ~ 7 planetary radii. This result also agrees with hydrodynamic model

simulations for a H-rich early Earth thermosphere by Tian *et al.* (2005). An idea that EUV exposed upper atmospheres expand dynamically to larger distances is also confirmed by more sophisticated models such as that of Chassefière (1996), and Tian *et al.* (2008a, b, 2009) which matched subsonic outflow solutions to Jeans boundary conditions at the exobase.

Figure 5 shows the thermospheric temperature profiles between the mesopause and exobase altitudes for a H-rich Earth-like planet which is exposed to 1 (present day), 2.5 and 5 times the present solar EUV flux value. One should keep in mind that we presented evidence in the previous sections that the EUV flux of the young Sun was much larger than 5 times that of the present Sun during the first 700 Myr after the Sun arrived at the ZAMS (see Table 7). During such stellar activity periods hydrogen-rich terrestrial planets may appear more as small gas giants with upper atmospheres which expand several tens of planetary radii.

5.2 Response of N₂-rich atmospheres to high EUV fluxes

Tian *et al.* (2008a, b) developed an advanced multi-component, hydrodynamic thermosphere model to study self-consistently the response of the Earth's thermosphere under EUV flux values from 1 (present) to 20 times that of the present-days solar value by assuming a present-time Earth atmospheric composition. Their model is validated against observations (MSIS-00: Hedin, 1991) and in agreement with other thermospheric models (e.g., Roble *et al.*, 1987; Smithtro and Sojka, 2005a, b) applied to the present day Earth. Figure 6 shows the thermospheric temperature profiles as a function of solar/stellar EUV flux values up to the corresponding exobase levels for various EUV cases. The results of these studies indicate that the thermosphere of Earth or a similar exoplanet with a present day Earth-like atmospheric composition should have switched from hydrostatic regime I to the non-hydrostatic regime II, if the thermospheric temperatures exceeded a critical value of about

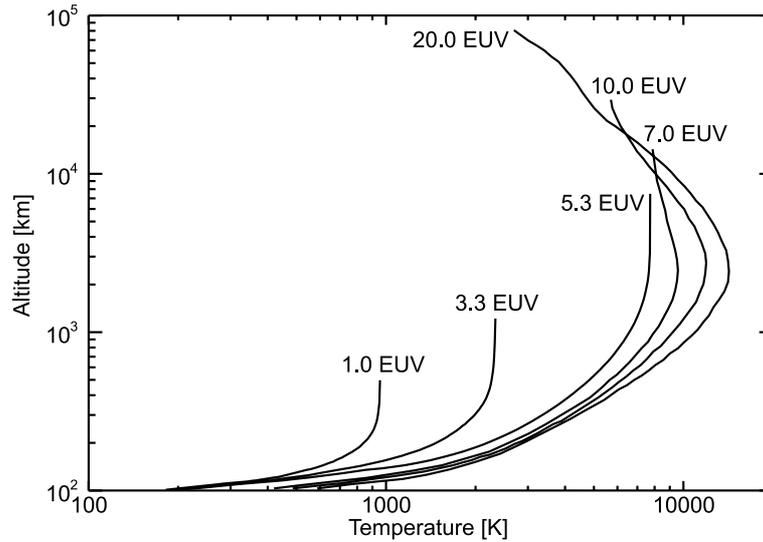


Fig. 6. Thermospheric temperature profiles between 100 km and the corresponding exobase levels for 1 (present), 3.3, 5.3, 7, 10 and 20 times higher EUV solar fluxes than today, applied to Earth with the present time atmospheric composition (Tian *et al.*, 2008a; Lammer *et al.*, 2008). One can see that the thermosphere switches in a non-hydrostatic regime for $T \geq 8000$ K and expands dynamically up to distances beyond 10 planetary radii.

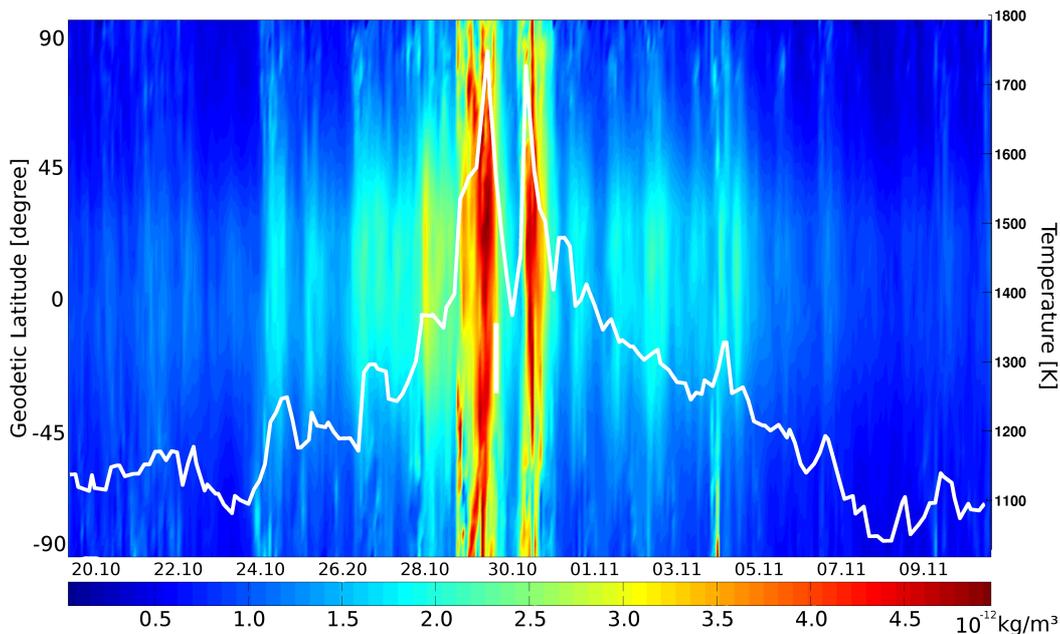


Fig. 7. Density impact of a X17.2 flare (Oct. 2003) observed by the GRACE satellite and temperature variation at orbit height is represented by the white curve (measurements taken at a local time of 16:19).

8000 K. For the Earth, these critical temperatures correspond to solar/stellar EUV flux values ≥ 6 times that of the present Sun.

One can also see that the exobase distance which separates the collision dominated thermosphere from the collisionless exosphere moves from ~ 500 km up to ~ 12.7 Earth-radii above the surface for a 20 EUV (4.2 Gyr ago) solar flux which would move about ~ 3.7 Earth-radii above the present-day average subsolar magnetopause location of ~ 9 Earth-radii. Under such extreme SXR and EUV conditions the magnetosphere would not protect the upper atmosphere, the planet experiences Venus or Mars-like solar

wind-atmospheric plasma interaction and as a result the exosphere could experience strong non-thermal escape.

The conjecture that the upper atmosphere of the Earth responds strongly to the solar X-ray, SXR and EUV radiation and particle fluxes can also be seen and studied from density enhancements during extreme solar events which are observed by low Earth orbit satellites (e.g., Sutton *et al.*, 2006; Shematovich *et al.*, 2011). Figure 7 shows the atmospheric neutral densities determined from accelerometer measurements onboard the low Earth orbiter Gravity Recovery And Climate Experiment (GRACE) satellite, where the neutral densities are determined with a 60 second resolution at an

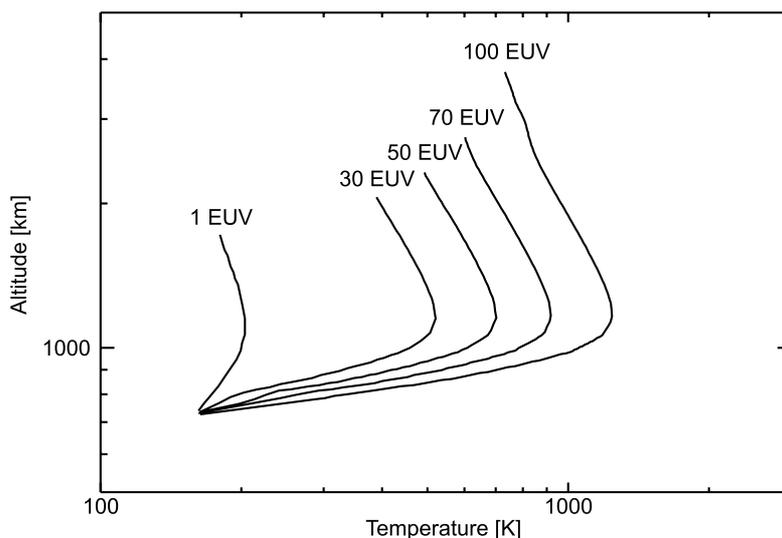


Fig. 8. Thermospheric temperature profiles between the mesopause and the corresponding exobase levels for 1 (present), 30, 50, 70, 100 times higher EUV solar fluxes than today, expected at Saturn’s large satellite Titan with the present time atmospheric composition.

orbital distance between 470 km at the perigee and 520 km at the apogee. The period under investigation covered ten days before and after the so-called Halloween geomagnetic storm (X17.2 flare) in 2003 (Sutton *et al.*, 2006). During this particular extreme solar event the upper atmosphere of the Earth was heated and as a result expanded with a density increase at the satellites orbit of $\approx 400\%$ compared to that at quiet conditions (Shematovich *et al.*, 2011). Related to that, also the temperature at the satellite position, calculated using the “Jacchia-Bowman model” (Bowman *et al.*, 2008), increases significantly from ~ 900 K up to ~ 1800 K. By applying a test-particle model to the expanded thermospheres given in Tian *et al.* (2008a) and a young solar wind which was about 100 times denser and twice as fast as today one obtains N^+ ion pick up loss rates in the orders of $\sim 10^{29}$ – 10^{30} s^{-1} (Lichtenegger *et al.*, 2010). Due to the extended thermosphere-exosphere region the Earth’s entire current N_2 atmosphere (~ 700 mbar) could have been stripped within a few million years.

In view of large uncertainties in the solar wind evolution discussed in Section 3.3 Lichtenegger *et al.* (2010) estimated also the ion pick up loss rate of these N_2 atmospheres by assuming present-day solar wind parameters. Although the loss rates are about a factor 10 lower, Earth would have lost its present nitrogen atmosphere within some tens of million years at the latest when the EUV flux was > 10 times from that of the present Sun. The main reason for these large escape rates is the EUV heated and expanded large solar wind plasma interaction area which is not protected beyond a present Earth-type magnetospheric shield.

Because the Earth’s nitrogen inventory is not lost, Lichtenegger *et al.* (2010) argued that these results indicate that the CO_2 amount in the early N_2 -rich terrestrial atmosphere was at least two orders of magnitude higher than the present-time level. Such an amount was needed to confine the upper atmosphere after the onset of the geodynamo within the shielding magnetosphere and thus might have protected it from complete destruction.

However, the model simulation results show that planets with the size and mass \leq of the Earth with even a 95% CO_2 atmosphere would experience problems when the EUV flux reaches values which are ≥ 30 times that of the present Earth (Tian, 2009; Lichtenegger *et al.*, 2010). Moreover, photochemically produced low energy (“hot”) atoms and additional thermospheric heating due to energetic solar/stellar particles could have additionally enhanced the atmospheric escape. On the other hand the expected initial CO_2 reservoir on early Earth was perhaps much larger compared to its nitrogen reservoir, so that a non-thermal loss to space of several bars of CO_2 during the first 100–300 Myr would have only slightly modified the CO_2 budget of the Earth.

Figure 8 shows the expansion of the second N_2 -type atmosphere of a solar system body, namely Saturn’s large satellite Titan. As pointed out by Lammer *et al.* (2008) self consistent hydrodynamic models applied to Titan’s thermosphere would be needed to study the photochemistry changes due to the different solar flux in the past to examine the modification in thermospheric temperature. Such studies are important for investigating to which distance the exobase could expand due to EUV heating by the young Sun and if Titan’s upper atmosphere could have reached classical hydrodynamic “blow-off” conditions or not.

5.3 Response of CO_2 -rich atmospheres to high EUV fluxes

High amounts of CO_2 , like on present Venus, Mars, and most likely on early Earth before 4 Gyr ago can cool the thermosphere much better than relatively dry Earth-like N_2/O_2 -type atmospheres, so that the exobase level remains closer to the planet’s surface and its upper atmosphere is better protected against erosion by the solar or stellar wind. For the more moderate EUV-flux period non-thermally photochemically produced suprathemal “hot” O atoms populate the exospheres of Venus and Mars. Gröller *et al.* (2010) developed a sophisticated Monte Carlo model and applied it recently to Venus during low and high solar EUV activity. One can see from Fig. 9 that the hot O atom corona densi-

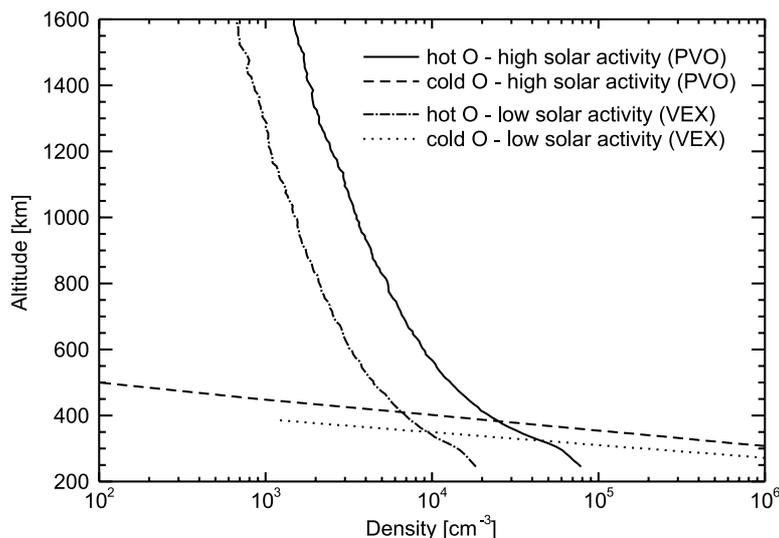


Fig. 9. Photochemically produced supra-thermal “hot” O atom density profiles at present Venus during low and high solar activity conditions compared with the bulk atmosphere.

ties for low solar activity (VEX solar conditions) are about a factor of 2–3 smaller than for high solar activity. Non-thermal hot O atoms populate the environment around the planet up to large distances where they can interact with the solar wind. This behavior indicates that higher EUV fluxes during the young Sun periods should have also resulted in denser hot O coronae around Venus and Mars, resulting in higher pick up loss rates.

Kulikov *et al.* (2006) studied the expected O^+ ion pick up loss rates over Venus’ history by using EUV data from solar proxies discussed in Section 2 and a range of solar wind plasma densities and velocities which were expected for the young active Sun. These authors used a diffusive-gravitational equilibrium and thermal balance model, which was applied for a study of the heating of the early thermosphere by photodissociation and ionization processes, exothermic chemical reactions, and cooling by CO_2 IR emission in the $15 \mu m$ band. The results of their study yield expanded thermospheres with exobase altitudes between about 200 km for present EUV flux values and about 1700 km for 100 times higher EUV fluxes after the Sun arrived at the ZAMS. The exobase temperatures yield values of about 8000 K for the earliest EUV fluxes after the Sun arrived at the ZAMS.

However, one can expect that the high EUV flux will dissociate CO_2 molecules in the thermosphere which results in less IR-cooling and in a possible non-hydrostatic upper atmosphere environment with high thermal escape rates for dissociated C and O atoms. By considering the uncertainties discussed on the solar wind parameters expected for the young Sun as discussed in Section 3.3, with the model discussed in Kulikov *et al.* (2006) one obtains loss rate from a CO_2 -rich early Venus atmosphere which is equivalent to about 2–10 bars during the planets’ history. From the present knowledge of the stellar wind parameters discussed in this work we expect that the high mass loss cases shown in Kulikov *et al.* (2006), where Venus could have lost much more atmosphere than 10 bar over its lifetime may be over-

estimations based on solar wind parameters which are too large.

If Venus evaporated its initial H_2O inventory during that period, the planets upper atmosphere would have been hydrogen-rich (Chassefière, 1996) as discussed in Section 5.1 and a dense hydrogen corona would have protected the initial N_2 and CO_2 inventory from escape until the evaporated water inventory was lost by thermal and non-thermal escape processes.

Tian *et al.* (2009) and Tian (2009) applied an advanced 1-D multi-component, hydrodynamic, planetary thermosphere-ionosphere model and a coupled electron transport-energy deposition model to Mars and CO_2 -rich atmospheres of “super-Earths”. With this model they could self-consistently calculate the ionization and excitation rates of thermospheric species together with the connected electron heating rates and related EUV heating efficiencies. It was found that due to the high EUV fluxes (>10 times that of the present Sun) CO_2 molecules can be dissociated so that less thermospheric IR-coolers are available which resulted in higher heating efficiencies and exobase temperatures compared to previous model simulations (Lammer *et al.*, 2006; Kulikov *et al.*, 2006). When extrapolating the model to early Mars, Tian *et al.* (2009) found that the exobase expands to ~ 900 km altitude and the exobase temperature reaches ~ 800 K for EUV flux values which are about 10 times the present solar value. As one can see from Fig. 10, for higher EUV fluxes such as 20 times the present solar value the exobase moved to distances of ≥ 10000 km with exobase temperatures >2500 K. One main reason for the nonlinear response of the upper atmosphere to the EUV flux is that the main IR cooling agents such as CO_2 , are destroyed by the high SXR and EUV fluxes, so that thermospheric cooling becomes less efficient.

According to their results the thermal neutral atom C and O escape fluxes $>10^{11} \text{ cm s}^{-1}$ from early Mars were as high as the outgassing fluxes which depended on ini-

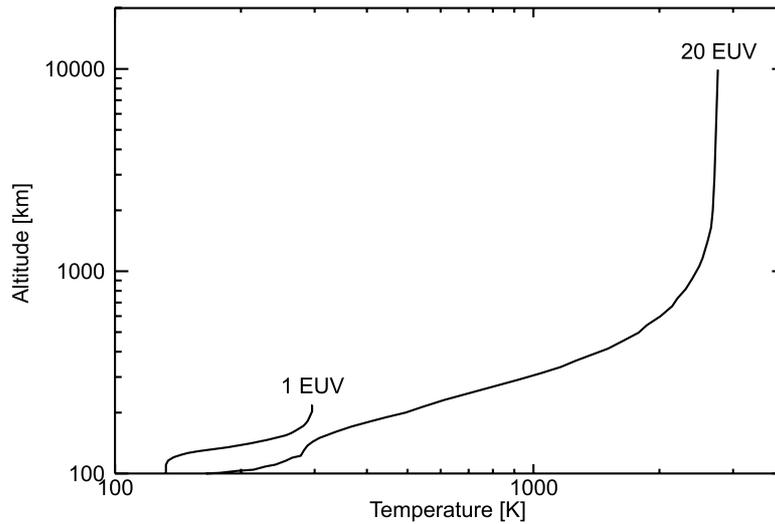


Fig. 10. Thermospheric temperature profiles for present day Mars between 100 km and the corresponding exobase levels for present time solar EUV flux values (1) (Fox and Hać, 2009) and the martian thermosphere when it is exposed to a 20 times higher EUV flux (~ 4 Gyr ago according to Tian *et al.* (2009)). The thermosphere temperature reaches ~ 3000 K and expands dynamically upward to distances of about 3 planetary radii.

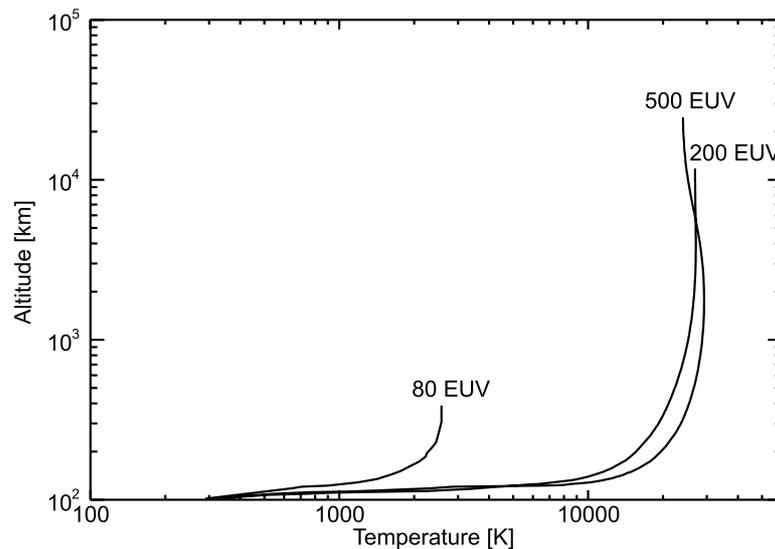


Fig. 11. Thermospheric temperature profiles according to Tian (2009) for “super-Earths” with ~ 6 Earth masses and CO_2 -type atmospheres within orbits of M-star habitable zones between 100 km and the corresponding exobase levels when their thermospheres exposed to a 80, 200 and 800 times higher EUV flux.

tial CO_2 inventories (e.g. Phillips *et al.*, 2001) during the early Noachian when the EUV flux was even larger than 20 times of the solar value. Due to the EUV heating and expansion of the upper atmosphere an early martian magnetic field could have not most likely protect the ion escape of C^+ and O^+ atoms which originated from dissociation of CO_2 molecules. By applying the same test-particle model, as in the study of early Earth by Lichtenegger *et al.* (2010), to the upper atmosphere profiles and estimated martian magnetopause distances as given in Dehant *et al.* (2007), we can show that the expanded thermosphere-exosphere can interact with the solar wind plasma which results in pick up C^+ ion escape fluxes of $> 10^{11} \text{ cm s}^{-1}$ before 4 Gyr ago.

Thus, due to the high activity of young Sun and the low gravity of Mars a dense early Noachian martian atmosphere

may not have built up during the first 300–400 Myr after the planet’s origin. A question remains when the outgassing rate was larger compared to the thermal and non-thermal neutral escape rates and what role impacts play in the formation or removal of the atmosphere (Pham *et al.*, 2009). These processes may have played a role at the end of the Noachian so that a CO_2 atmosphere with a surface pressure of a few 100 mbars and a resulting warmer climate may have been established so that Mars warmed up sporadically ~ 3.5 –4 Gyr ago. However, such an accumulated CO_2 atmosphere should have been lost since about 4 Gyr ago by solar wind erosion (ion pick up, cold ion outflow), sputtering and photochemical hot atom escape after the martian magnetic dynamo stopped working.

Tian (2009) applied the same model to more massive ~ 6

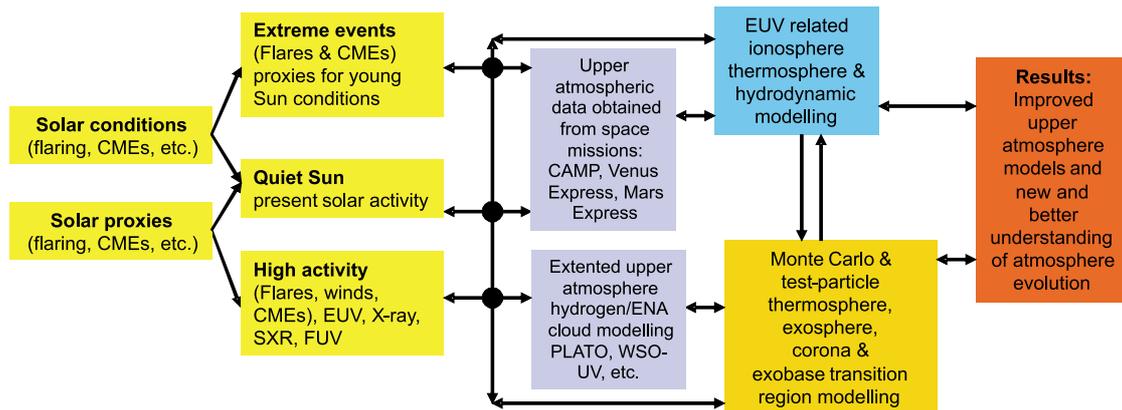


Fig. 12. Illustration chart about the general approach and connections between observations and modeling efforts which should lead to the development of state-of-the-art thermosphere-exosphere models and to a better understanding of aeronomically related atmospheric evolution aspects for terrestrial planets.

Earth-mass hypothetical CO₂-rich exoplanets. The results indicate that such atmospheres are stable against thermal escape, even for M stars' EUV flux enhancements as large as 1000 compared to the present Sun. On the other hand one can see from Fig. 11 that the highly irradiated thermospheres of such planets may also expand to altitudes beyond 10000 km above the planets surface. These results are in agreement with previous studies of Khodachenko *et al.* (2007) and Lammer *et al.* (2007) who estimated the atmospheric erosion of CO₂-rich exoplanets due to CME-induced ion pick up within close-in habitable zones of active M-type dwarf stars. They found that Earth-mass CO₂-rich exoplanets that have no, or weak, magnetic moments may lose tens to hundreds of bars of atmospheric pressure, or even be stripped of their whole atmospheres due to CME-induced O⁺ ion pick up at orbital distances ~ 0.2 AU. Although the formation of magnetospheres and possible locations of a magnetopause are not well known for such planets, expanded atmospheres of “super-Earths” could have been subject to stellar wind erosion. However, one should note that previous studies focused either on thermal effects and escape (e.g., Tian *et al.*, 2008a, 2009; Tian, 2009) or on non-thermal processes (e.g., Kulikov *et al.*, 2006, 2007; Terada *et al.*, 2009; Lichtenegger *et al.*, 2010) although both processes are strongly coupled and should not be separately studied. As illustrated in Fig. 12, to understand really how planetary atmospheres can “survive” the extreme X-ray, SXR and EUV activity of their host stars after they arrived at the ZAMS, one has to combine thermal and non-thermal upper atmosphere processes by developing a self-consistent multi-species ionospheric-thermospheric exosphere model which couples thermal and non-thermal processes, for studying hydrostatic and non-hydrostatic dynamically expanding thermosphere-exosphere regions, including the investigation of the production of hot atoms, their collisional interaction and transport within highly EUV exposed thermosphere-exosphere regions, and their contribution to the thermospheric heat budget.

6. Testing Atmosphere Evolution Scenarios via UV-transit Observations of Earth-like Exoplanets Orbiting Dwarf Stars

As discussed in the previous section, dynamically expanded thermospheres and related coronae should form if SXR and EUV fluxes ≥ 5 EUV and light atoms such as hydrogen populate the upper atmosphere, or if no IR-cooling molecules are present. As a consequence the upper atmosphere expands beyond a possible magnetospheric shield and hydrogen energetic neutral atoms (ENAs) are produced via charge exchange collisions between the atmosphere and charged solar/stellar wind particles (Lammer *et al.*, 2011a, b). As discussed recently by Holmström *et al.* (2008), Ekenbäck *et al.* (2010) and Lammer *et al.* (2011a) ENAs have been observed around all Solar System planets when a spacecraft was equipped with an accurate instrument.

Lammer *et al.* (2011b) pointed out that, the observation and theoretical modeling of hydrogen and ENA-clouds around a terrestrial exoplanet would indicate that its thermosphere is not in hydrostatic equilibrium but experiences hydrodynamic expansion beyond its magnetospheric shield. Therefore, by observing planetary hydrogen clouds and related ENA-coronae around transiting terrestrial exoplanets in UV (Lyman- α attenuation) we will learn more on how the early Solar System planetary atmospheres have evolved. Because transiting Earth-size exoplanets and their hydrogen clouds are too small relative to the size of Sun-like G-stars to detect them, Lammer *et al.* (2011b) proposed to study hydrogen clouds around such planets within orbits of M dwarf stars. Due to the large number of M-type stars in our neighborhood (Scalo *et al.*, 2007) and their small size and comparably long periods of strong stellar activity in X-rays, SXR, and EUV M dwarf stars may represent the most promising candidates for the detection of hydrogen clouds and related ENAs.

Terrestrial planets such as the Earth are formed via collisions between planetary embryos and by accumulation of planetesimals in the presence of the primordial solar/stellar nebula gases (e.g. Alibert *et al.*, 2010; references therein). At this early formation stage young terrestrial planets are surrounded by massive primordial gas envelopes which are

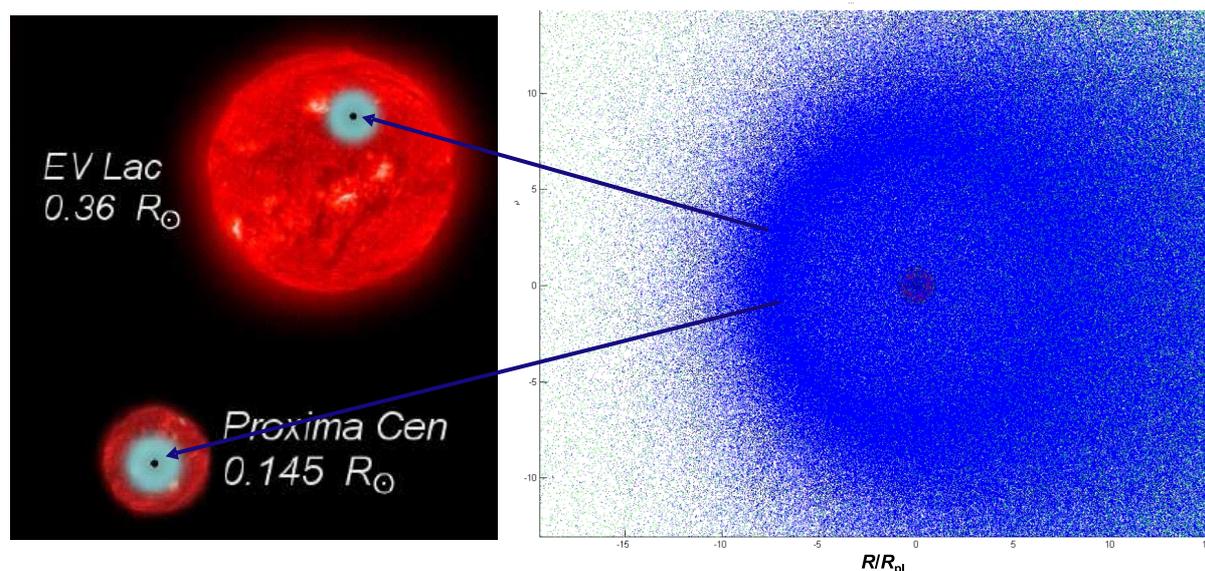


Fig. 13. Left side illustrates the size relation of exoplanets around the two M-stars EV Lacertae and Proxima Centauri with a transiting Earth analogue with an extended hydrogen atmosphere of $7R_{\text{Earth}}$. Right side shows a preliminary model simulation of an extended hydrogen corona and stellar wind produced hydrogen ENAs around an Earth-size planet with an expanded hydrogen-rich thermosphere of $7R_{\text{Earth}}$ (Lammer *et al.*, 2011b).

mainly composed of hydrogen and helium. According to Lammer *et al.* (2011b) such examples may be:

- dense hydrogen envelopes which remained from the primordial nebulae;
- hydrogen-rich thermospheres produced via dissociation of CH_4 , NH_3 and H_2O molecules;
- ocean planets or planets which evaporate their initial H_2O reservoir;
- evaporation of H_2O oceans due to cometary or asteroidal impacts.

During this early stage in planetary evolution, heavy atmospheric species such as CO_2 , N_2 , etc. are most likely protected from escape to space until the dense primordial hydrogen envelope which populates the non-hydrostatic thermosphere is lost. The removal time of such hydrogen envelopes depends on the amount of available hydrogen, the SXR, EUV flux and plasma environment of the young host star. If this hypothesis is right, one can expect that hydrogen-rich upper atmospheres may be a common scenario at young terrestrial planets during tens or hundreds of Myr after their origin.

If we expect that similar processes produce hydrogen-rich thermospheres around many young terrestrial exoplanets, these bodies would produce huge planetary hydrogen coronae and related ENA clouds as observed around hot gas giants such as HD 209458b (Vidal-Madjar *et al.*, 2003; Holmström *et al.*, 2008; Ekenbäck *et al.*, 2010) and HD 189733b (Lecavelier des Etangs *et al.*, 2010). Figure 13 shows an illustration of such transiting exoplanets and a modeled hydrogen cloud around an Earth-size test planet according to Lammer *et al.* (2011b). As one can see from Fig. 13 the produced hydrogen atoms populate a wide area around the planetary environment and a huge fraction of the host star should be covered with the outward flowing plan-

etary hydrogen atoms and produced ENA-cloud during a transit.

Since the past years the search for Earth-like exoplanets around low-mass M-type stars has become one of the major activities in stellar astrophysics. Besides several ground-based search projects such as CRILES, CARMENES, or HZPF (e.g., Lammer *et al.*, 2011b; and references therein) ESA's M3 candidate PLATO (*PLANetary Transits and Oscillations of stars*) will also be dedicated to detect transiting terrestrial exoplanets within and outside of habitable zones, and determine their basic parameters with unprecedented accuracy (Catala and the PLATO team, 2009a, b). The PLATO mission has been selected for a definition study phase in the context of ESA's Cosmic Vision 2015–2025 program as an M-class mission with foreseen launch in 2017–2024 (Stankov *et al.*, 2010). The instrument concept of PLATO (Catala and the PLATO team, 2009a) consists of a multi-telescope approach. In total 32 refractive telescopes (pupil size about 120 mm each) are combined to produce highly photometrically accurate lightcurves covering a wide dynamic range. In addition, two so-called fast telescopes with fixed color filters and higher read-out cadence are used for seismic studies of the brightest targets. PLATO is planned to be launched into an orbit around the L2 Lagrange point. Two target fields will be monitored for at least 2–3 years with high duty cycle. In a “step-and-stare” phase several fields will be monitored for shorter durations (2–5 months each). In total, PLATO will cover about 50% of the sky.

If selected, this mission will follow the ongoing CoRoT and Kepler missions to obtain a new sample of transiting exoplanets, focussing on terrestrial planets around bright stars ($m_V \leq 11$ mag). This sample will not only provide planetary systems with highly accurately known basic planetary and stellar parameters, but will also provide access

to a well-constrained age of planetary systems. Planetary radii will be determined from PLATO transit observations. Combined with a world-wide ground-based follow-up campaign to obtain radial velocity measurements, planet masses will be derived. However, previous transit searches have shown that well constrained stellar parameters are crucial to derive accurate planet parameters. PLATO therefore combines photometric transit search with astroseismology of the central stars of planetary systems. These astroseismological studies are further keys to obtain a well constrained age (goal: 10% accuracy; Kjeldsen *et al.*, 2009) of stars, and therefore planetary systems.

The search will include planets of all orbital periods and with all physical sizes, with focus on bright central stars of solar type. PLATO will perform a full characterization of the observed very bright stars using seismic analysis. The highest priority sample, therefore, are about 20 000 dwarfs and sub-giants with $m_V \leq 11$ mag for which Earth-sized planets can be detected and accurate seismic analysis is performed. A smaller sample of very bright stars ($m_V \leq 8$ mag) will provide targets for future spectral characterization of the detected planets. Finally, a large sample of stars (more than 245 000 cool dwarfs/subgiants with $11 \text{ mag} < m_V < 13$ mag) will be monitored to detect terrestrial planets, but without detailed astroseismology. What is interesting in the context of our work outlined in this section, PLATO will also search for exoplanets around more than 10 000 nearby M-type dwarfs (down to $m_V = 15\text{--}16$ mag), including planets within their habitable zones.

After the discovery of Earth-size exoplanets with PLATO, follow-up observations with the the upcoming World Space Observatory-UltraViolet (WSO-UV) mission could be carried out. The WSO-UV is an international space mission devoted to UV spectroscopy and imaging which is included in the Federal Space Program of the Russian Federation with a launch date around 2014–2015 (Shustov *et al.*, 2009; Lammer *et al.*, 2010). The observatory includes a 170 cm aperture telescope capable of high-resolution and long slit low-resolution spectroscopy, and deep UV and optical imaging and is designed for observations in the UV-range (Shustov *et al.*, 2009). The WSO-UV space observatory will be equipped with instruments which can carry out:

- High resolution spectroscopy of point sources in a wavelength range 1020–3200 Å by means of 2 high resolution echelle set ups with a sensitivity which is about 10 times better than that of the HST/STIS.
- Long-slit resolution spectroscopy in the 1020–3200 Å range.
- High resolution imaging in the 1500–2800 Å range with a field view of $1' \times 1'$.
- High resolution imaging in the 2000–7000 Å range with a sensitivity which is about 3 times higher than the HST/ACS system.
- Molecular hydrogen mapping in extended, faint astrophysical objects with a large field far UV (1150–1900 Å) camera.

The WSO-UV's core program will focus also on exoplanetary atmospheres and astrochemistry in the presence of

strong UV radiation fields (Gómez de Castro *et al.*, 2009; Shustov *et al.*, 2009). In Lammer *et al.* (2011b) it was shown that it may be possible to observe transits in UV of Earth-size exoplanets with extended hydrogen coronae if they orbit around small stars.

Thus, transit follow-up observations of exoplanets discovered by ESA's currently studied PLATO mission in the UV-range of terrestrial exoplanets around M-type stars with space observatories such as the WSO-UV would provide a unique opportunity to shed more light on the early evolution of Earth-like planets, including those of our own Solar System. Moreover, as pointed out in Lammer *et al.* (2011a) observations of exoplanets with various ages, together with hydrogen cloud and ENA modeling efforts could be used for the reconstruction of the solar wind history as well as testing magnetic dynamo evolution scenarios of gas giants.

7. Conclusions

By analyzing solar and astrophysical observations and phenomena we present strong evidence that the Sun and similar stars underwent much more active phases after they arrived at the ZAMS. The enhanced activity reveals itself in the form of strong high energy radiation emissions, frequent flares, and a denser and faster stellar wind. Results from the “Sun in Time” project suggest that the coronal X-ray and EUV emissions of the very young main-sequence Sun were ~ 1550 and ~ 80 times stronger than those of the present Sun, respectively. We showed that such an energy and particle environment certainly had an impact on the origin and evolution of Solar System planets and their atmospheres. Due to the heating of the much higher solar/stellar SXR and EUV flux the planetary thermospheres and exobase levels extended to higher altitudes than at present time, which resulted in larger solar wind-atmosphere interaction areas, no protection of intrinsic magnetospheres, the production of huge hydrogen ENA-clouds and higher non-thermal escape rates. Under certain activity conditions of the young Sun or planetary host star hydrostatic equilibrium could not have been kept. Finally we addressed the possibility of studying the detection and follow-up observations of extended upper atmospheres around exoplanets by near future space observatories such as ESAs PLATO mission in combination with the WSO-UV space telescope. Such observations would enhance our knowledge regarding the interaction of the host star's plasma environment and with Earth-like exoplanets itself, as well as shed light into the evolutionary stage of their atmospheres.

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